

L02-Planck's law: quantization

Saturday, August 27, 2016 13:39

- * Motivation for [re]studying Planck's law
 - it was the origin of quantum mechanics
 - practice with a distribution (Boltzmann)
 - review of classical (E&M) waves:
 - dispersion relation & quantization (modes in a cavity)
 - introduce quantization of classical waves $\rightarrow \hbar$

* Blackbody history (Gasiotowicz 1-1, wiki: Planck's Law)

- 1858 Balfour Stewart - Lamp black absorbs most light
- 1859 Kirchhoff - universality of emissivity / absorption
 - introduced Blackbody radiation - noted importance
- 1865 Tyndall - observed peak in spectrum
- 1880, 81-6 Crowl, Langley - better data 3-d data
- 1894 Wien - $f(\lambda T) / \lambda^5$ based on entropy argument
- 1896 Wien - $e^{-c/\lambda T} / \lambda^5$ high freq. cutoff
- 1898 Lummer & Kurlbaum high quality data
 - to distinguish exact nature of Planck's law
- 1899 Lummer & Pringsheim: $\sim \lambda^{-5} e^{-c/\lambda T}$
- 1900 Rayleigh 'heuristic formula: $C_1 T \lambda^{-4} e^{-c_2/\lambda T}$
- 1900 Planck: $2hc^2/\lambda^5 (e^{hc/\lambda T} - 1)^{-1}$ empirical $C \lambda^{-5} / e^{c/\lambda T} - 1$
 - combined Rayleigh + Wien using entropy
- 1905 Rayleigh-Jeans law: $\frac{8\pi}{\lambda^4} kT$ or $\frac{8\pi \nu^2}{c^3} kT$
- 1911 Ehrenfest - "UV catastrophe" Wien-Nobel prize
- 1918 Planck accepts physical quantization - Nobel prize.

- * energy density vs. intensity
(same as vacuum conductivity through an aperture)
imagine particles in 4π , what is the flux through a hole?

$$\begin{aligned} d\Phi &= \vec{J} \cdot d\vec{a} = \rho \vec{v} \cdot d\vec{a} = \rho \int \frac{v d\Omega}{4\pi} \underbrace{\hat{v}(\theta, \phi) \cdot d\vec{a}}_{\cos\theta} \\ &= \frac{\rho v}{4\pi} \int d\phi \sin\theta d\theta \cdot \cos\theta = \frac{\rho v}{4\pi} \cdot 2\pi \int_0^1 u du = \frac{1}{4} \rho v \end{aligned}$$

$\frac{1}{2}$ particles to the right, $\frac{1}{2}$ from integration of $\cos\theta$

- * switching variables in a distribution:

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$$u_\nu(\nu) d\nu = u_\lambda(\lambda) d\lambda$$

$$c = \nu \lambda \quad \frac{d\lambda}{d\nu} = -\frac{c}{\nu^2}$$

$$u_\nu(\nu) = u_\lambda(\lambda) \left| \frac{d\lambda}{d\nu} \right| = u_\lambda\left(\frac{c}{\nu}\right) \cdot \frac{c}{\nu^2} \quad (-): \nu \text{ increases as } \lambda \text{ decreases.}$$

* Wien's law: Wien, Rayleigh-Jeans, Planck distributions

$$u_\lambda(\lambda) = \lambda^{-5} f(\lambda T) \quad \frac{hc}{\lambda kT} = x = \frac{h\nu}{kT} \quad u_\nu(\nu) = \frac{\nu^3}{c^4} f\left(\frac{c}{\nu} T\right)$$

$$\text{Wien: } f(\lambda T) = 8\pi hc \cdot e^{-x}$$

good as $\lambda \rightarrow 0$

$$\text{Rayleigh-Jeans: } f(\lambda T) = 8\pi hc \cdot \frac{1}{x}$$

good as $\lambda \rightarrow \infty$,

physically motivated.

$$\text{Planck: } f(\lambda T) = 8\pi hc \cdot \frac{1}{e^x - 1}$$

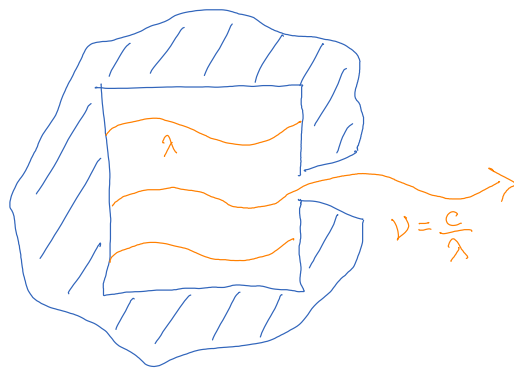
good everywhere

* Rayleigh/Jeans were able to explain the low-wavelength tail of the black body spectrum using basic physical principles:

Let, for simplicity, the black body be a cubic cavity with a small hole for the radiation to escape.

They assumed 1) the radiation was present in "modes" inside the cavity which satisfied boundary conditions.

2) the energy of each mode was random, and thermally distributed



The spectral energy density, defined as $u(\nu) = \frac{\text{energy}}{\text{volume} \cdot \Delta \text{frequency}}$ is given by the product: $u(\nu) = g(\nu) \cdot \bar{\epsilon}(\nu)$

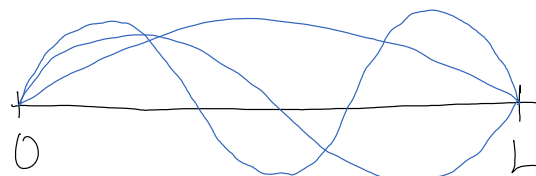
of modes between ν and $\nu + d\nu$ * average energy of each mode

Assumption #1: modes in the cavity:

This is a property of classical waves, but we will apply the same idea to quantum wave functions next class.

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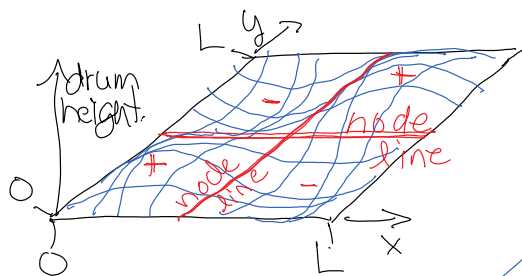
- 1) Consider a 1-d string:
clamped at $x=0$ & $x=L$



There are n half-wavelengths

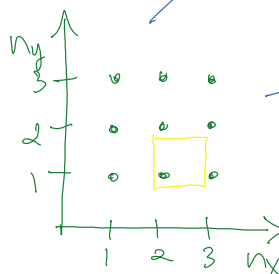
$$L = n \frac{\lambda}{2} \quad \text{or since } c = \lambda \nu \quad \nu = \frac{c}{\lambda} = \frac{cn}{2L}$$

- 2) Now consider a 2-d square drumhead. The modes will oscillate between evenly spaced node lines:

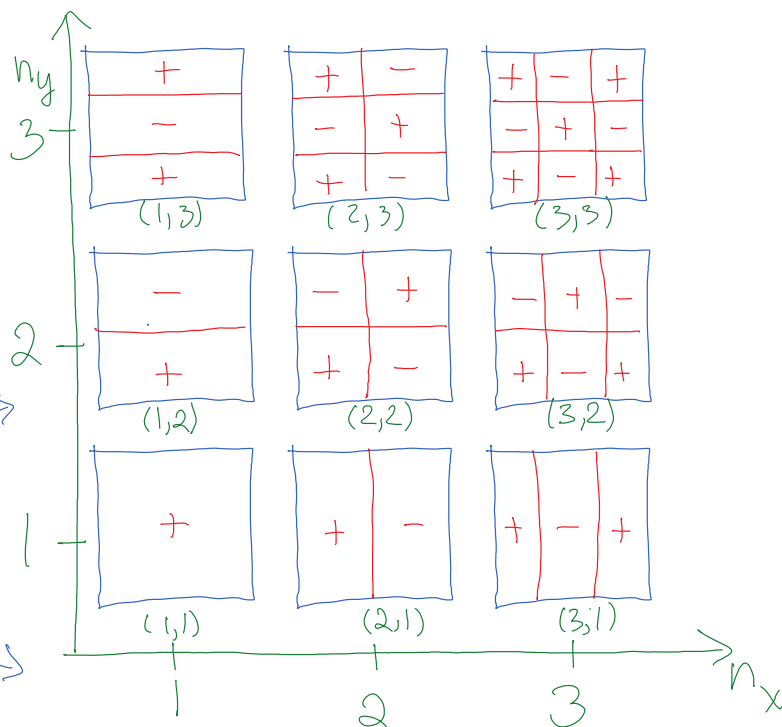


(2,2)-mode

Each dot represents one mode of vibration.

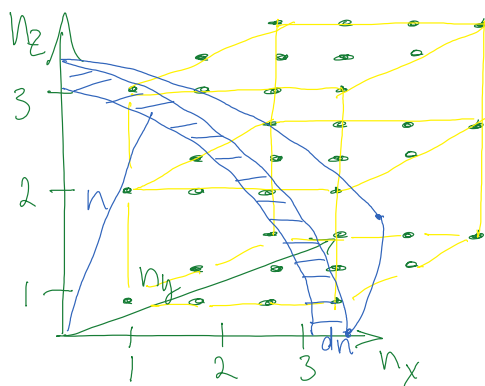


Note each dot occupies one unit of area in the (n_x, n_y) space.



- 3) In our 3-d cavity, modes have three sets of node planes. They are labelled (n_x, n_y, n_z) and have frequency $\nu = \frac{cn}{2L}$ where $n^2 = n_x^2 + n_y^2 + n_z^2$

The # of modes between ν & $\nu + d\nu$ is the volume of a shell of radius n and thickness dn



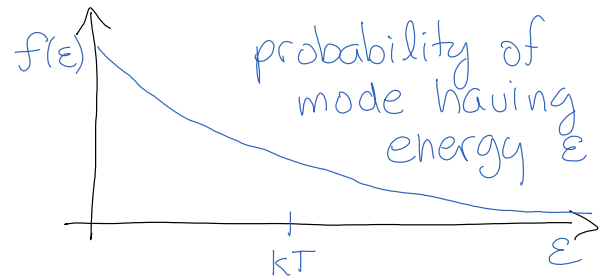
$$\int g(\nu) d\nu = \frac{1}{4} 4\pi n^2 \cdot dn = \frac{\pi}{2} \left(\frac{2L}{c}\right)^3 \nu^2 d\nu \quad \times 2 \text{ polarizations}$$

$$\left[g(\nu) d\nu = \frac{1}{8} 4\pi n^2 \cdot d\mathbf{n} = \frac{\pi}{2} \left(\frac{2L}{c}\right)^3 \nu^2 d\nu \right] \times 2 \text{ polarizations of light}$$

Note the number of modes is much larger at high ν .

Assumption #2: The thermal energy of each mode follows the Boltzmann distribution. $f_B(\epsilon) = e^{-\epsilon/kT}$

Lower energy states are exponentially more likely. The distribution broadens at higher temperature to accommodate more energy



To calculate $\bar{\epsilon}(\nu)$ we must take a weighted average:

$$\bar{\epsilon}(\nu) = \frac{\int \epsilon f_B(\epsilon) d\epsilon}{\int f_B(\epsilon) d\epsilon}$$

The denominator (normalization) is called the "Partition function"

$$Z(\beta) = \int_0^\infty e^{-\beta\epsilon} d\epsilon = \left. \frac{-1}{\beta} e^{-\beta\epsilon} \right|_0^\infty = \frac{1}{\beta} \quad \text{where } \beta = \frac{1}{kT}$$

There is trick to get the numerator from Z :

$$\int \epsilon e^{-\epsilon/kT} d\epsilon = \int -\frac{\partial}{\partial \beta} e^{-\beta\epsilon} d\epsilon = -\frac{\partial}{\partial \beta} Z(\beta) = \frac{1}{\beta^2}$$

$$\text{Thus } \bar{\epsilon}(\nu) = -\frac{\partial}{\partial \beta} \ln Z = \frac{-\frac{\partial Z}{\partial \beta}}{Z(\beta)} = \frac{1/\beta^2}{1/\beta} = \frac{1}{\beta} = kT \quad \text{independent of } \nu$$

This is the equipartition, which states the average energy per degree of freedom (D.O.F.) is $\frac{1}{2} kT$. The two D.O.F.'s are electric & magnetic fields.

$$\text{Thus } u_\nu(\nu) = \frac{8\pi\nu^2}{c^3} kT = \frac{8\pi h\nu^3}{c^3} \cdot \frac{1}{x} \quad \text{where } x = \frac{h\nu}{kT} = \frac{hc}{\lambda kT}$$

* How can we restore the missing UV cutoff at high frequency, and prevent infinite energy?

Planck was able to explain his fit to the data

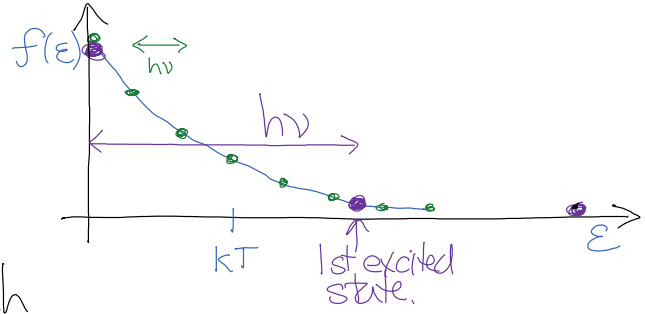
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Assumption #3: energy is limited to discrete packets or "quanta" which get larger with frequency $\boxed{\epsilon = h\nu}$

Instead of integrating $f(\epsilon)$, he summed over steps $\epsilon_n = nh\nu$

For low frequency packets, $\epsilon = h\nu$, the sum is close to the previous integral.



For high frequency packets $\epsilon = h\nu$, there is not enough thermal energy to excite even the first excited state, and thus the average energy of the mode is 0.

$$Z(\beta) = \sum_{n=0}^{\infty} e^{-\epsilon_n/kT} = \sum_{n=0}^{\infty} e^{-\frac{nh\nu}{kT}} = \sum_{n=0}^{\infty} a^n = \frac{1}{1-a} \quad a = e^{-\beta h\nu} = e^{-x}$$

$$\sum_{n=0}^{\infty} \epsilon_n e^{-\epsilon_n/kT} = -\frac{\partial}{\partial \beta} Z(\beta) = \frac{-1}{(1-a)^2} \cdot e^{-\beta h\nu} \cdot (-)h\nu$$

$$\text{so } \bar{\epsilon}(\nu) = \frac{\sum_{n=0}^{\infty} \epsilon_n f(\epsilon_n)}{\sum_{n=0}^{\infty} f(\epsilon_n)} = \frac{-\frac{\partial}{\partial \beta} Z(\beta)}{Z(\beta)} = \frac{h\nu e^{-x}}{1-e^{-x}} = \frac{h\nu}{e^x - 1} = \frac{kT \cdot x}{e^x - 1}$$

$$\text{and } u(\nu) = \frac{8\pi h\nu^3}{c^3} \cdot \frac{1}{e^x - 1} \quad \text{as indicated by experiment.}$$

The modern interpretation of Planck's radical assumption is that E&M waves come in packets of energy, "photons" & Thus waves exhibit particle-like behaviour.

This interpretation was cemented by Einstein's explanation of the photoelectric effect, and the explanation of Compton scattering in terms of point-particle kinematics.